

Enhancing the Vector Vector Vortex Coronagraph:

Achromatic design and chromatic filtering
Mitigation of susceptibility to aberrations and stellar size
Immunization to central obscuration, segment gaps and spiders

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Disclaimers:

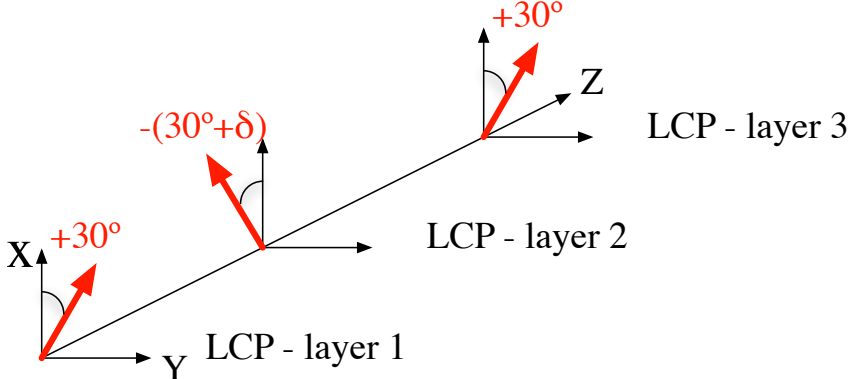
- Looking for ground-based L-band coronagraphy ? Go see Absil's poster and Delacroix's talk.
- If you don't know what the VVC is, go see my other poster first.

Achromatic 3-layer design

Based on the Pancharatnam phase achromatization principle:

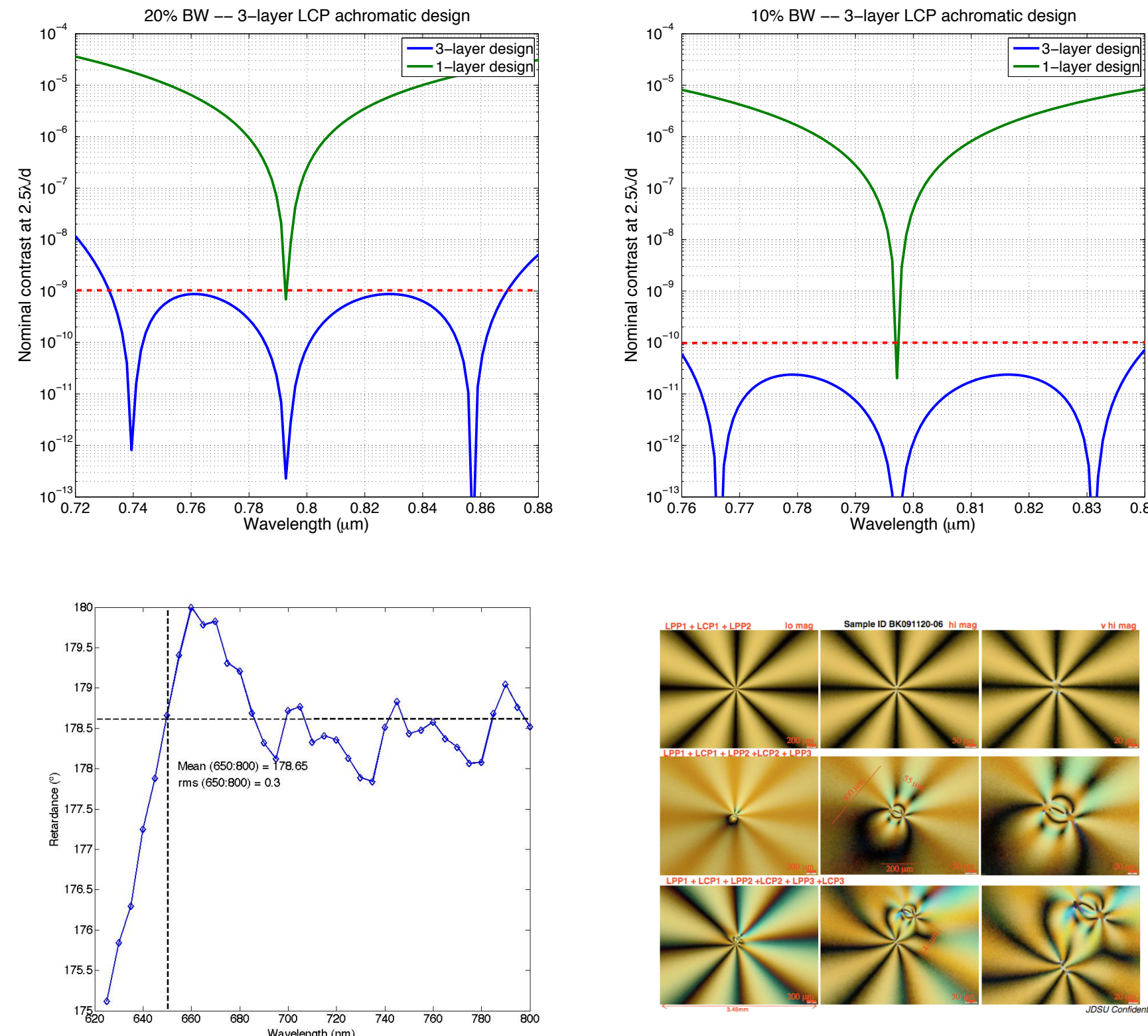
Used for decades to manufacture achromatic waveplates:

- Needs 3 identical LCP layers.
- Optical axis offset between the 3 layers.
- Net effect is an achromatic 180° retardance.



First attempt at realizing the 3-layer design (JDSU):

- Achromaticity good for 10⁻⁹ contrast.
- The current apparatus misses a layer registration mechanism.
- Central region messed up.
- New manufacturer, new technique (projection lithography vs contact lithography).



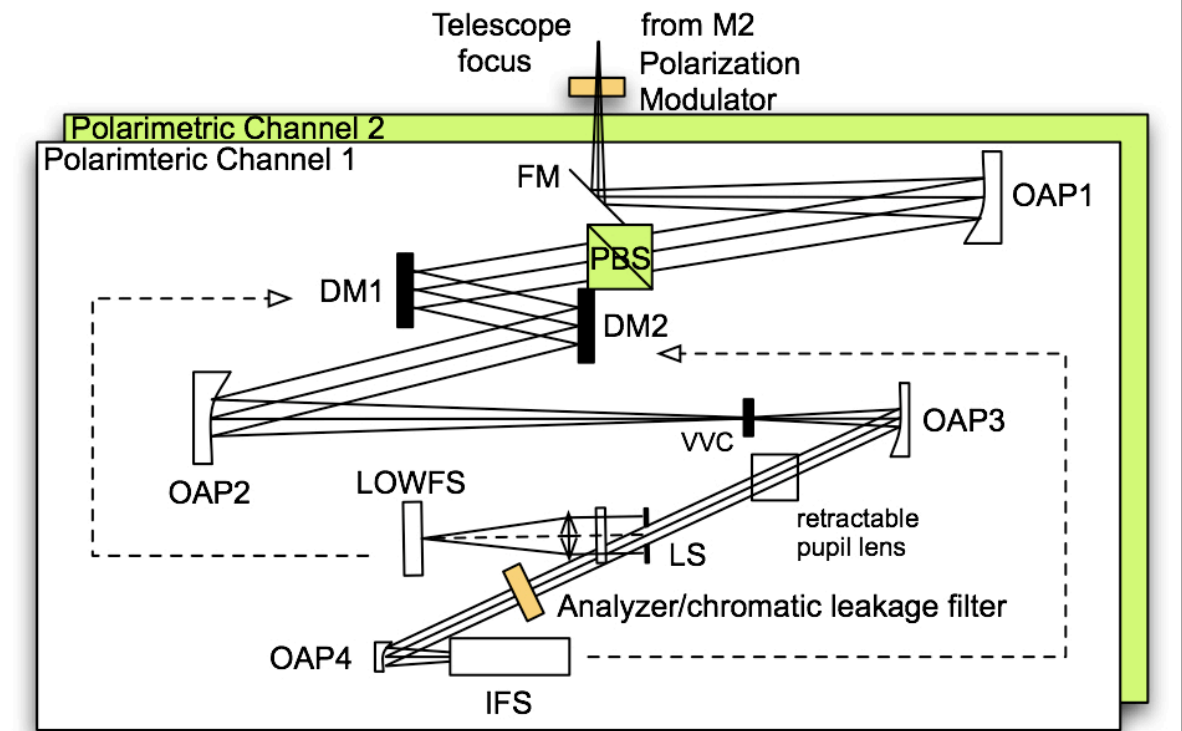
Residual chromatic leakage filtering - polarimetric differential imaging

$$J_v(r, \theta) = V \begin{bmatrix} e^{il\theta} \\ e^{-il_p\theta} \end{bmatrix} + L \begin{bmatrix} 1 & 0 \\ 0 & 1 \end{bmatrix}$$

Take advantage of the vectorial nature of the VVC:

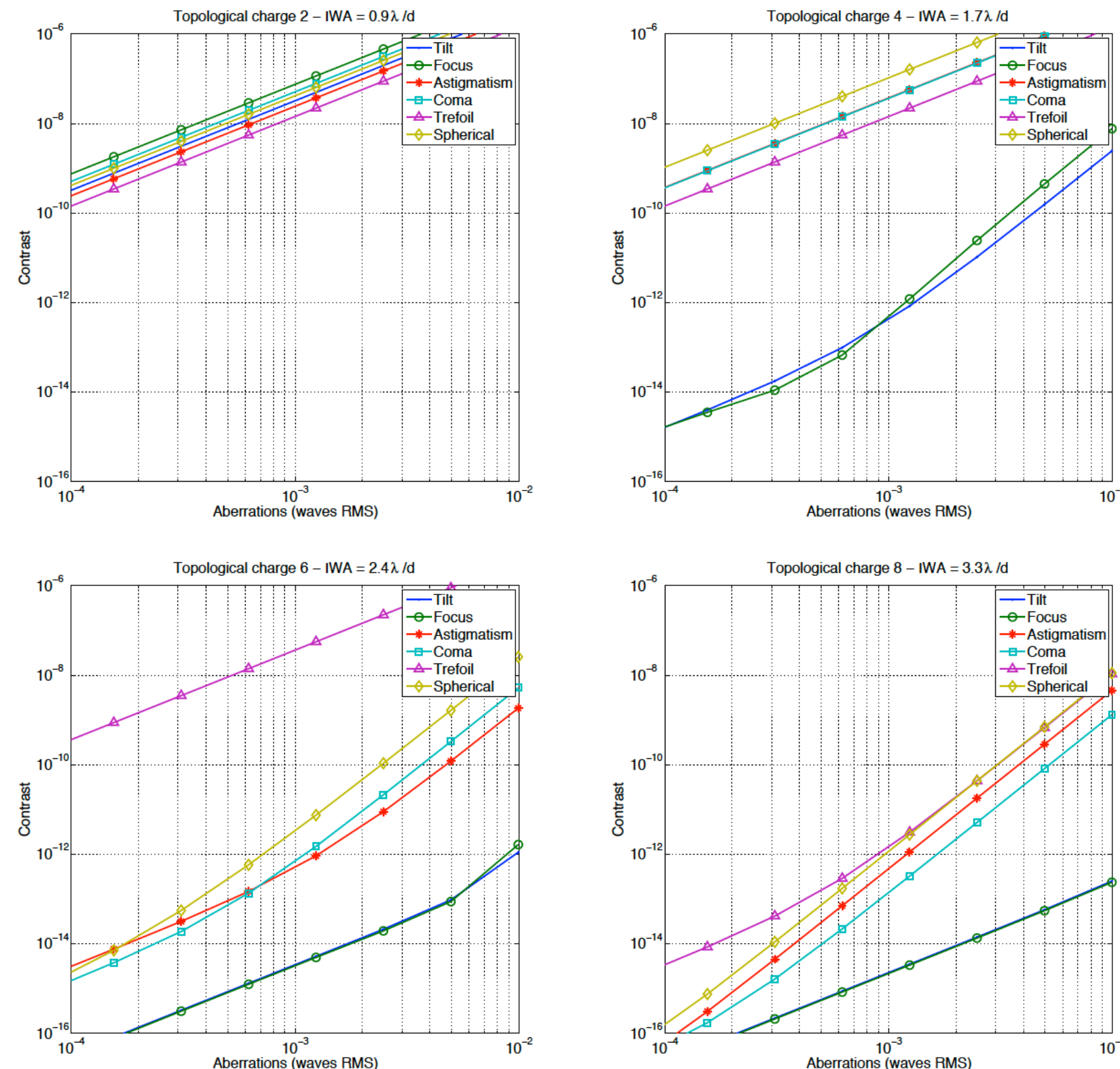
- The chromatic leakage has an orthogonal polarization to the main vortex term and any off-axis companion.
- Can be removed in polarized light by a simple analyzer.
- The gain in contrast of the filtering process is related to the quality of the analyzer as follows:

$$G \approx \left[\frac{1}{ER} + \Delta\phi_{QWP}^2 \right]^{-1}$$



Bandwidth	Modifier angle	Contrast at ~ 2.5 λ/d	Leakage filtering
10%	0.1	~ 10 ⁻¹⁰	< 10 ⁻¹⁰
20%	0.4	~ 10 ⁻⁹	< 10 ⁻¹⁰
40%	1.0	~ 10 ⁻⁷	~ 10 ⁻¹⁰
60%	2.5	~ 10 ⁻⁶	~ 10 ⁻⁹

Low-order aberration sensitivity as a function of the topological charge

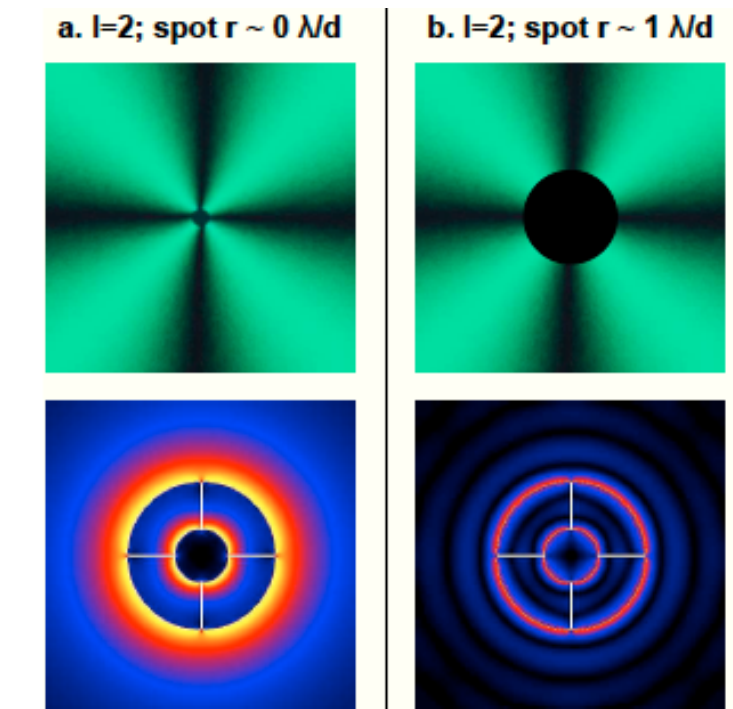


Sensitivity decreases with the topological charge, so does the IWA => Trade-OFF

Immunization to Central Obscuration

Solution 1. trade throughput with contrast:

- Hybridization of the VVC.
- Joint optimization of the central opaque spot and lyot stop size.
- Typically for 15-25% central obscuration:
 - Contrast ~ 10⁻⁷
 - Throughput loss of ~ 50-90 %
- See Mawet et al. 2010 (SPIE).



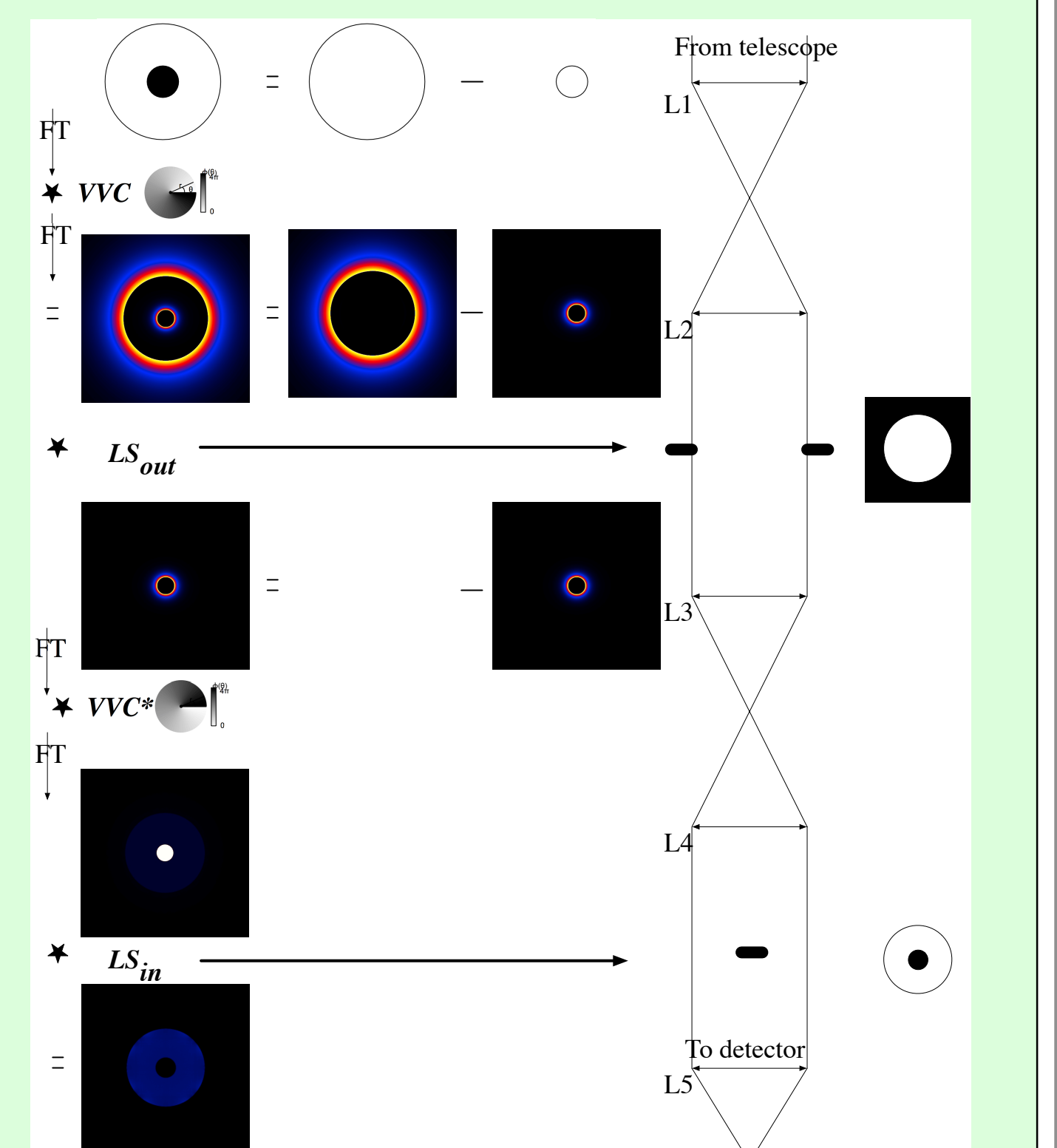
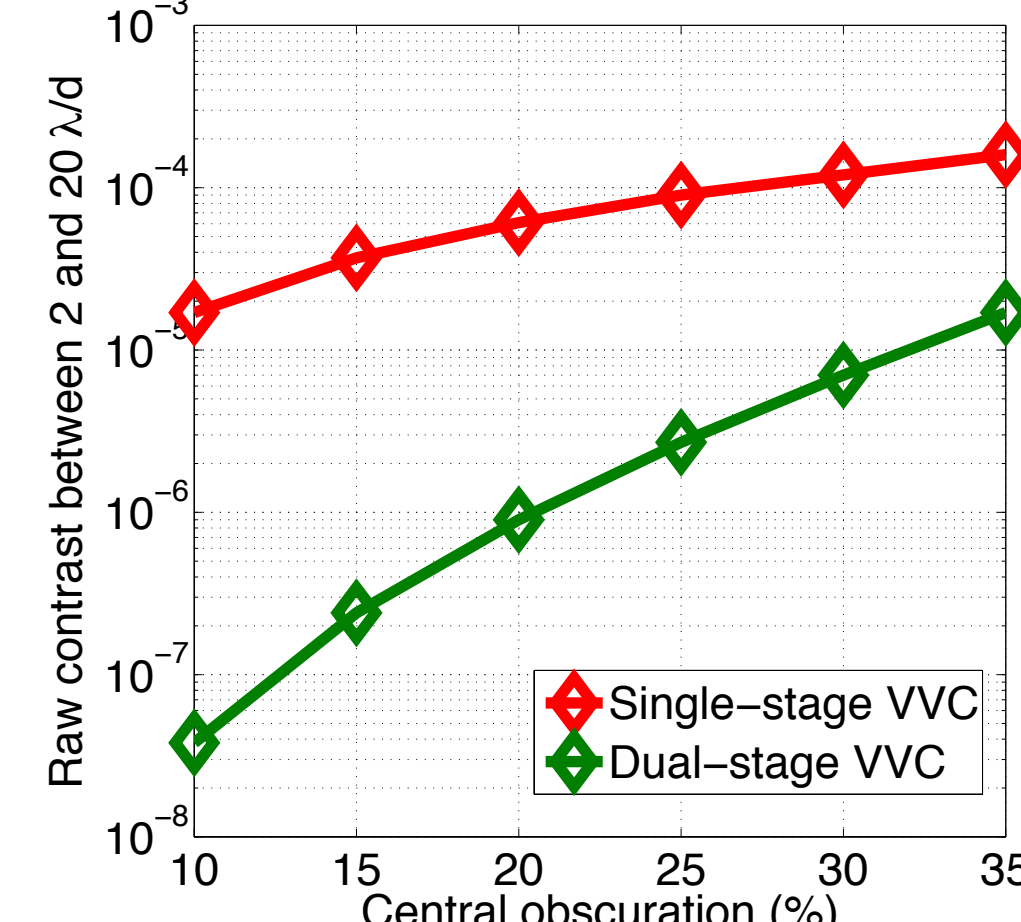
Hybrid mask: VVC + central spot, note that the spot is never bigger than the IWA

Pupil diffraction pattern (same scale)

Solution 2. the return of the jedi dual-stage coronagraph:

- Amazing property of the VVC: phase-induced amplitude redistribution.
- No THROUGHPUT LOSS !!!
- Known advantages of the multi-stage approach:
 - Reduced chromaticity.
 - Reduced sensitivity to stellar size/low-order aberrations.

Contrast vs central obscuration



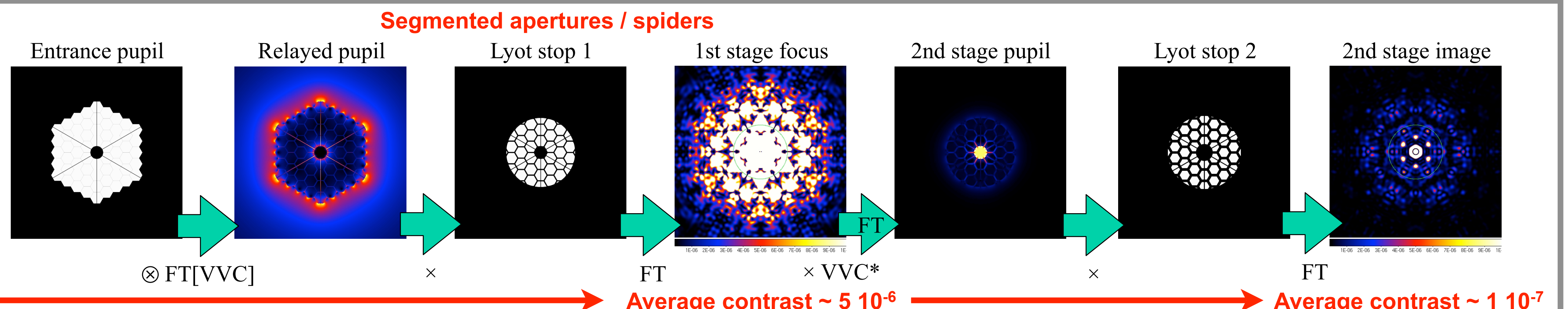
Stellar size matters

Contrast at 3λ/d as a function of star apparent diameter and topological charge

Topological charge	IWA (λ/d)	s = 0.01λ/d	s = 0.1λ/d
2	0.9	1 × 10 ⁻⁸	1 × 10 ⁻⁶
4	1.7	3 × 10 ⁻¹³	3 × 10 ⁻⁹
6	2.4	8 × 10 ⁻¹⁸	8 × 10 ⁻¹²
8	3.3	2 × 10 ⁻²²	1 × 10 ⁻¹⁴

The VVC is not afraid of segmented pupils, and spiders !

- Use the edge-enhancement property of the VVC to our advantage.
- Careful optimization of the Lyot stop allows mitigating the high spatial frequency content of spiders and gaps between segments.
- The two-stage VVC provides additional leverage to remove diffracted light.
- Conclusions: deep raw contrast of 10⁻⁷ are not ruled out on segmented obscured pupils with carefully designed coronagraphs !!!



Perspectives

The VVC coronagraph family provides a set of handles to mitigate chromaticity, sensitivity to low-order aberrations, stellar size, central obscuration, gaps between segments and spiders. A particular property of the VVC, called phase-induced amplitude redistribution can be used in a dual-stage layout to dramatically reduce the effect of central obscuration while maximizing the throughput, reduce chromaticity, stellar leakage and sensitivity to low-order aberrations. **The results presented here and recent advances in wavefront control technology and post-processing techniques put phase-mask coronagraphs on the table again as potential solutions for extremely large and large ground-based telescopes.**